Imaging Pipeline Software

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Math705 Research Project

Introduction

An imaging pipeline for Image Synthesis is designed to take data either gathered by radio interferometry telescopes, or generated to simulate those telescopes, and form an image of an area of the sky. This software will be similar to that which will be used for the Square Kilometer Array, a project that AUT is involved in, where the same processes will have to take place.

This project will follow a Design Science methodology where software will be the generated artefact used to experimentally investigate image synthesis. The project will start with the gathering of knowledge on the techniques used in an imaging pipeline. It will also involve gaining knowledge on solutions to complications such as concurrency control and the mapping of visibilities to a grid. Then the project will involve implementing the techniques in the form of a Java program and then will be tested using visibility input data available in the High Performance Computing Research Laboratory. The software will be testing against other pipelines and changes will be made to try to improve its performance.

The output from the developed software, using visibility data as input, will be an image of the sky. The images produced by the pipeline will be analyzed to compare it against the known sky images for the data sets to validate whether the techniques are implemented properly and potentially look at its performance.

It is expected that an imaging pipeline will be developed with the capability for image synthesis. Also expected to gain knowledge in the three main steps involved in the pipeline, namely gridding, (inverse) Fourier transform, and deconvolution, as well as some techniques for algorithm optimization.

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Literature Review

## Synthesis Imaging

The resolution of radio telescopes can be increased by using pairs of telescopes (baselines) and taking the product of the received signals. This resolution can be changed by increases the separation of the baseline, rather then increasing the size of the individual telescopes. This method popularized by the work of (Ryle & Hewish, 1960) states that using these baselines it produces “exactly the same result as that obtained by using the complete large aperture”. This technique allowed for cheaper production of much larger apertures and the eventual development of the techniques used now.

These techniques gather Fourier domain data in the form of a visibility, however the way in which they are sampled is non-uniform, so we must place it on a rectangular grid. This process is known as gridding and the methods used now are based on the work by (Brouw, 1975). These visibilities V(u, v) fall upon the plane in which the baselines are setup, for a wider coverage of this place more baselines can be added and could also be moved around.

With more modern telescopes being developed, moving them around became a substantial task and instead the rotation of the earth can be used to move these points around the plane. An image of these points on the V(u, v) plane can be seen on the left of Figure 1.

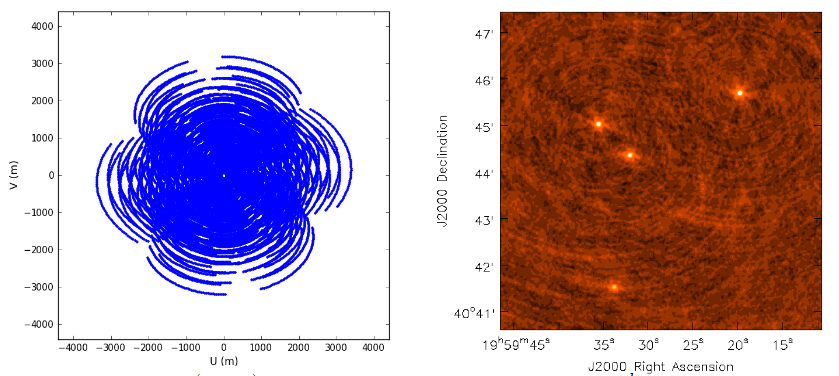
As there are gaps within the V(u, v) plane the image is a “dirty image”, this can be seen on the right of Figure 1.

Figure 1. V(u,v) plane showing data points. (Rau, 2012) Retrieved from Australian Telescope National Facility, from https://www.atnf.csiro.au/research/radio- school/2012/lectures/tue/RVU\_ImagingDeconvolution.pdf

## Fast Fourier Transform (FFT)

A Fourier Transform is a process for signal-processing and analysis. (Brigham, 1988) states that the extent of the use the process is as follows “biomedical engineering, imaging, analysis of stock market data, spectroscopy, metallurgical analysis, nonlinear systems analysis, mechanical analysis, geophysical analysis, simulation, music synthesis”. It is widely regarded as one of the most important algorithms based on its impact in so many areas. Simply put a Fourier Transform is used to show different parts of a continuous signal, however for Interferometry an Inverse FFT (iFFT) is used as we are combining the amplitude and phase of the signal to form an image. For the performance of the pipeline an inverse Fast Fourier Transform will be used. Using such a method is based upon the work of (Hogg, MacDonald, Conway, & Wade, 1969). The algorithm used was first discovered by Gauss and later rediscovered by (Cooley & Tukey, 1965) which also notes that “Wherever possible the use of N = r^m with r = 2 or 4 offers important advantages” which impacts the design of the pipeline.

## Gridding

Gridding is the process of mapping the data collected onto a rectangular grid so that it may be processed by the iFFT and then displayed as an image. Early techniques for places the visibility data on a grid involved placing the visibilities upon the closest grid point that aligns with the plane on which they were gathered and either adding then all together or averaging them out. Early methods where used by (Hogg, MacDonald, Conway, & Wade, 1969). However, this led to artifacts forming and therefore a limited application for the process. An alternative method was first used by (Brouw, 1975) and would take a weighted value based the distance between local grid point and the point of the visibility. By designating a “support” area around the local grid point the data can be added to these areas. An ideal gridding method was given by (O'Sullivan, 1985) with his gridding algorithm that used a sinc function was “computationally efficient” and resulted in “arbitrarily small artifact levels”, however this function would give infinite extent to the support. This is not ideal computationally for the gridder so instead the use of convolution kernels with a set support are commonly used.

A simplified approach to gridding follows these steps. For every visibility, find the closest grid point to the data on the V(u, v) plane, using the convolution kernel the data point is spread across the support region.

## Deconvolution

Once an image is formed from the Fourier Transform is it called a ‘dirty image’, as seen in Figure 1, this is due to the effects of having limited sampling of the V(u, v) plane. The process of Deconvolution can be used to ‘CLEAN’ the image. This method by (Högbom, 1974) uses the original V(u, v) data to form a “dirty beam”, then by taking away the dirty beam from the points of the dirty image with the greatest brightness you are left with a residual image. By iteratively doing this process a clean image is left behind.

This was expanded upon by (Clark, 1980) to make it more efficient. His method involves using more FFT’s in a major and minor cycle to subtract points away from the dirty image.

Further examples of the ‘CLEAN’ algorithm and the Maximum Entropy technique by (Skilling & Bryan, 1984) are compared by (Cornwell & Bridle, 1996).

Methods

## Design of pipeline

The pipeline will consist of three main sections, the Gridder, the inverse Fourier Transform, and the Deconvolution. Before that I will break down some design choices made before coding began.

For the language to be used I decided on Java, this is due to it being the language I am most familiar with. Java is the language primarily used in my computer science papers at AUT. It is also more then capable of handling the size of the data I will be using and supports parallelization.

In order to test my pipeline, I will use a dataset generated by the HPCRL at AUT. By using this data I can better check how well the Fourier Transform and deconvolution is working as it comes with a perfect image made from the real components of the sky. This perfect image comes from the data being generated, with captured data we never know the true sky image and so it is les useful to compare. As well as a perfect image the data contains Visibilities.csv, a file containing ~23000 data points. Each visibility consists of its locations along the (u, v) plane, following by its value as a complex number. The dataset also includes the ProlateSphre

Results

Discussion of Results

Conclusion

**References**

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Appendix